Constraining the neutrino emission of gravitationally lensed Flat-Spectrum Radio Quasars with ANTARES data

The ANTARES collaboration

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Abstract. This paper proposes to exploit gravitational lensing effects to improve the sensitivity of neutrino telescopes to the intrinsic neutrino emission of distant blazar populations. This strategy is illustrated with a search for cosmic neutrinos in the direction of four distant and gravitationally lensed Flat-Spectrum Radio Quasars. The magnification factor is estimated for each system assuming a singular isothermal profile for the lens. Based on data collected from 2007 to 2012 by the ANTARES neutrino telescope, the strongest constraint is obtained from the lensed quasar B0218+357, providing a limit on the total neutrino luminosity of this source of $1.08 \times 10^{46}\,\mathrm{erg\,s^{-1}}$. This limit is about one order of magnitude lower than those previously obtained in the ANTARES standard point source searches with non-lensed Flat-Spectrum Radio Quasars.

Keywords: gravitational lensing, neutrino astronomy

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Contents

1 Introduction

Active Galactic Nuclei (AGNs) are longstanding candidate sources for the highest-energy cosmic rays, the origin of which is still unknown. The blazar subclass is comprised of BL Lac objects and the more luminous Flat-Spectrum Radio Quasars (FSRQs), the relativistic jets of which are pointed at a small angle to the line of sight [1]. The matter content of AGN jets is still heavily debated, as both leptonic and hadronic models can in principle explain the broad band spectral energy distribution of blazars [2]. As of now, for electromagnetic radiation there is no simple observable that allows us to distinguish between both types of models. The detection of neutrinos from such jets might therefore help solve this long-standing issue.

In hadronic models for AGN jets, neutrinos are predicted to be emitted along with gamma-rays by processes involving the interaction of accelerated hadrons with the surrounding medium and radiation fields, and the subsequent production and decay of pions and kaons (see e.g. [3, 4] and references therein).The neutrino spectrum should therefore be closely related to the hadron spectrum, usually assumed to be a power-law with spectral index $\gamma \simeq -2$ as suggested by the theory of diffuse shock acceleration [5, 6]. The IceCube neutrino telescope has recently provided evidence for a cosmic component in the diffuse high-energy neutrino flux [7], part of which could originate from a population of unresolved extragalactic sources, possibly blazars [8]. Yet, no individual neutrino source has so far been identified, and, given the low expected neutrino fluxes, the published upper limits do not allow different jet models to be distinguished. BLLacs are interesting targets to derive generic constraints on the jet emission mechanisms as they can be found at relatively small redshifts. For blazar populations which are typically distributed at much larger distances, such as the FSRQs, we argue that gravitational lensing can help enhance the sensitivity to neutrino emission.

Gravitational lensing of electromagnetic radiation from distant astrophysical sources is a well-known prediction of Einstein's theory of general relativity [9, 10]. Since the first detection of multiple images of a gravitationally lensed quasar, QSO 0957+561 [11, 12], hundreds of lens systems have been discovered and studied, opening new perspectives both in astrophysics and cosmology (see [13] for a recent review).

Due to their very low masses (m_ν do not exceed $\sim 1\,\mathrm{eV}$ [14]), neutrinos are expected to undergo the same phenomenon of gravitational lensing as photons. Possible configurations for neutrino lensing by massive astrophysical objects have been theoretically studied in the literature, e.g. [15–18]. For distant neutrino sources gravitationally lensed by an interposed galaxy (or galaxy cluster), the magnification of the neutrino flux is expected to be equal to that of the photon flux.

Here, we make use of this similarity between the lensing of photons and of neutrinos to significantly lower the upper limits on neutrino emission from FSRQs, by using the luminosity boost provided by gravitational lenses that magnify some distant radio-loud AGN. We illustrate this method by performing a search for an excess of neutrino events in the direction of distant lensed FSRQs in the field of view of the ANTARES neutrino telescope.

The remainder of this paper is organized as follows. Section 2 describes the ANTARES neutrino telescope, which is used to detect neutrino events from gravitationally lensed sources, followed in section 3 by a description of the candidate sample. We present and discuss our results in section 4. Specifically, we show that this strategy leads to an improvement in the upper limit on the intrinsic neutrino luminosity of FSRQs.

2 The ANTARES neutrino telescope and data sample

The ANTARES neutrino telescope [19] is located in the Mediterranean Sea, about $40\,\mathrm{km}$ offshore from Toulon, France ($42°48'$N, $6°10'$E). The detector consists of 885 photomultiplier tubes (PMTs) mounted on twelve vertical lines anchored to the seabed at a depth of 2475m, with a typical inter-line separation of $\sim 65\,\mathrm{m}$. The operation principle is based on the detection of Cherenkov light induced by relativistic muons produced in the charged-current interactions of high-energy neutrinos with matter surrounding the detector. The knowledge of the time and amplitude of the light pulses recorded by the PMTs allows the reconstruction of the trajectory of the muon and to infer the arrival direction of the incident neutrino. To limit the background from down-going atmospheric muons, the design of ANTARES is optimised for the detection of up-going muons produced by neutrinos which have traversed the Earth. Its instantaneous field of view is $\sim 2\pi$ sr for neutrino energies $E_\nu \gtrsim 100\,\mathrm{GeV}$. Further detail on the above can be found in refs. [19–21].

The ANTARES collaboration has developed several strategies to search for cosmic neutrino candidates, also in association with other cosmic messengers such as cosmic-, X-, and gamma-rays, and gravitational waves (e.g. [22–26]). A search for neutrino point-like sources was conducted with the data sample corresponding to the first four years of operation of the detector, 2007–2010 [27]. This search has recently been extended to the years 2011–2012, for a total detector livetime of 1338 days [28]. The selection criteria have been optimised to search for E^{-2} neutrino fluxes from point-like astrophysical sources. Upgoing events have been selected to reject the bulk of background atmospheric muons. Additional cuts on the quality of the muon track reconstruction have also been applied to eliminate events that correspond to downgoing atmospheric muons which are misreconstructed as upgoing. Most of the remaining events are atmospheric muon neutrinos which constitute the primary background for cosmic neutrino searches. The final sample contains 5516 neutrino candidates with a predicted atmospheric muon neutrino purity of around 90% and an estimated median angular resolution of 0.38° [28].

The above-mentioned sample was searched for an excess neutrino flux in the direction of 51 candidate neutrino sources (including the Galactic Centre). This list of sources was

established mainly on the basis of the intensity of their gamma-ray emission as observed by *Fermi* [29] and H.E.S.S. [30] and includes five non-lensed FSRQs. No statistically significant excess has been found in the direction of any of the candidate sources. The corresponding 90% confidence level (C.L.) upper limits on the neutrino flux, assuming an E^{-2} spectrum, have been derived; they are the most restrictive to date in a significant fraction of the Southern Sky. In the next section, we show how these limits can be improved upon by using lensed sources.

3 Search for neutrino emission from distant, lensed blazars

3.1 Determination of the magnification factor

The strategy presented here relies on an estimation of the magnification factor for each lensed system, as obtained from photon observations. To estimate the magnification of a given lensed image, a model is required for the mass distribution of the lens. Here, we utilise the simplest model able to account for the image morphology and magnification ratios in each lens system: the singular isothermal ellipsoid (SIE). SIE models are also a reasonable description of the mass distributions of elliptical galaxies (e.g., [31]). We adjust the parameters of SIEs with the "gravlens" modeling software [32] to fit image positions and flux ratios. We use image positions from infrared H-band HST observations from the CASTLES project,[1] which have uncertainties of 0.003 arcsec. We assume the centre of mass of each lens galaxy to be its centre of brightness. The centre for each SIE model is then the centre of brightness of each lens galaxy. The positions of the SIEs have uncertainties between 0.01 and 0.05 arcsec (CASTLES). Because the flux ratios may be influenced by microlensing, the SIEs cannot reproduce them to better than a 10% uncertainty, which we adopt for our flux uncertainties. The resulting uncertainties for the estimates of magnifications for the lensed images in our sample are about 15% (10% for quadruple-image systems), excluding systematic uncertainties.

3.2 Candidate distant lensed blazars

The most interesting lensed blazar candidates are B0218+357 [33] and PKS 1830–211 [34]: they are visible to ANTARES and have also been detected by *Fermi*. B0218+357 is a double-image lens system with an extended jet at redshift $z = 0.96$, while the intervening lensing spiral galaxy has $z = 0.68$ [35]. We obtain magnification values of 12.3 and 8.5 for the two images based on the SIE model. These estimates are consistent with the observed flux ratios at radio wavelengths [36]. B0218+357 is also the first lensed system for which a clear γ-ray measurement of a time delay for two images has been reported [37]. The measurement utilised flares detected with the *Fermi* Large Area Telescope (LAT). Although the LAT is unable to resolve the two images, the flares were of sufficient amplitude (peaking at ~ 20–50 times the quiescent flux) to conduct an autocorrelation analysis. The result was a delay of 11.46 ± 0.16 days, generally consistent with previous measurements at radio wavelengths. Some of the uneven structures of the flares implied that microlensing may be occurring in this system. To establish the effect of microlensing on the magnifications of the images that we use in this paper, models of the distribution of stars in the lens galaxy and their kinematics are necessary. Further analysis is underway to model the behavior of the γ-ray flares (Falco et al., in preparation).

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[1]http://www.cfa.harvard.edu/castles/.

Flares in γ-rays from *Fermi*-LAT were also reported for PKS 1830–211, a double-image lens system with a separation of ~ 1 arcsec [38]. This radio-loud blazar is at $z = 2.51$ and is gravitationally lensed by a spiral galaxy at $z = 0.89$. Based on the SIE model, the magnification values that we obtain for the two images are 3.7 and 1.5. A time delay of 27.1 ± 0.6 days between the images was estimated at radio wavelengths [39]. Evidence for a time delay in the *Fermi* LAT γ-ray data, consistent with the radio measurement, has also been reported for this source [40]. Observations with ALMA that overlapped with some of those of *Fermi* suggested microlensing at sub-mm wavelengths [41]. Their measurements also revealed chromatic variations in the flux densities of the images at these wavelengths, which may arise from microlensing or from intrinsic variability of the blazar jet. Our estimates for the magnifications of the images of PKS 1830–211 yield a magnification ratio of ~ 2.4, compared with the estimated magnification ratio of ~ 1.5 at radio wavelengths [39]. Because this source is highly variable (both extrinsically and intrinsically), the estimates of magnification ratios are affected and also variable. We adopt our model magnifications as fiducial values.

We also include in our study B1422+231 [42] and B1030+074 [43] which are two lensed blazars in the ANTARES field of view, although with no associated *Fermi* detection so far. B1422+231 is a four-image quasar at $z = 3.62$ lensed by a galaxy at $z = 0.34$. Its optical spectrum shows broad emission lines, and it is categorized as an FSRQ in the Multifrequency Catalogue of Blazars [44]. The magnification values of the four images as obtained from the SIE model are 16.0, 15.0, 11.1 and 0.9. B1030+074 is a two-image system with a variable source exhibiting jet structure at $z = 1.54$, and a lens galaxy at $z = 0.60$. It is a blazar of uncertain type in [44]. We obtain magnification values of 2.4 and 0.27 for the two images.

3.3 Search for neutrino emission

We search for neutrino emission from the above sources following the same procedure as described in [27, 28]. Neutrino events within $20°$ of the source are used to construct an unbinned likelihood function including both signal and atmospheric background components. For each source, this likelihood is then maximised with respect to the number of signal events. The sensitivity of the analysis is evaluated through the generation of large numbers of pseudo-experiments simulating background and signal. No significant excess of neutrinos above the expected background has been found for the four candidate sources. The corresponding 90% C.L. upper limits on the neutrino flux have been derived using the Feldman & Cousins approach [45]; they are given in table 1 along with the limits on the five non-lensed FSRQs obtained in the previous ANTARES analysis.

4 Results and discussion

From the upper limits on the neutrino flux at Earth $\phi_\nu^{90\text{CL}}$ (given in table 1), we derive limits on the intrinsic neutrino luminosities of the sources. The total isotropic emitted power in high-energy neutrinos ($\gtrsim 1\,\text{TeV}$) for a blazar at luminosity distance d_L is

$$L^{90} = \frac{1}{\mu}\, 4\pi d_L^2\, \phi_\nu^{90\text{CL}}, \tag{4.1}$$

where μ is the magnification factor ($\mu = 1$ for non-lensed sources). The corresponding values are given in table 1 for both lensed and non-lensed FSRQs considered in this study. Note that the value L^{90} itself is not a source-intrinsic quantity. It depends on the viewing angle via the Doppler factor. In lack of well-constrained viewing angles for most blazars,

Name	z	d_L	$\phi_\nu^{90\mathrm{CL}}$	μ (n)	L^{90}
3C 279	0.54	3.12	3.45	1	6.44
PKS 1454–354	1.40	10.2	1.70	1	33.8
3C 454.3	0.86	5.54	2.39	1	14.1
PKS 1502+106	1.84	14.3	2.31	1	90.5
PKS 0727–11	1.60	12.0	2.01	1	55.3
PKS 1830–211	2.51	21.0	1.89	5.20 (2)	30.8
B0218+357	0.96	6.35	2.91	20.8 (2)	1.08
B1422+231	3.62	32.7	2.46	43.0 (4)	11.7
B1030+074	1.54	11.5	2.26	2.67 (2)	21.5
QSO 0957+561	1.41	10.3	-	2.80 (2)	-

Table 1. Redshift z, luminosity distance d_L (in units Gpc), 90% C.L. flux limit (in units $10^{-8}\,\mathrm{GeV\,cm^{-2}\,s^{-1}}$), magnification factor μ and number of lensed images n (when relevant), and upper limit on intrinsic luminosity L^{90} (in units $10^{46}\,\mathrm{erg\,s^{-1}}$) for the FSRQs considered in this analysis.

however, conclusions about the intrinsic luminosities can be derived by applying statistical methods [46].

The results are summarised in figure 1 where the intrinsic luminosity limit is plotted against the luminosity distance of the source. For the lensed systems, the error bar account for the uncertainty in the determination of the magnification factor as discussed in section 3.1. We have added for comparison the luminosity limits obtained for the BLLacs in the ANTARES list of sources, most of which are located at a much closer distance [28]. The isotropic power limits obtained for the non-lensed FSRQs are between 6×10^{46} and $9 \times 10^{47}\,\mathrm{erg\,s^{-1}}$. These powers can be compared with the bolometric luminosities of AGNs, which are typically in the range 10^{44}–$10^{47}\,\mathrm{erg\,s^{-1}}$ [2, 47, 48], but can rise up to $10^{49}\,\mathrm{erg\,s^{-1}}$ for some hadronic jet models such as in the synchrotron proton blazar interpretation of 3C 279 [49].

One directly sees from figure 1 that the limits derived from lensed FSRQs are stronger than those corresponding to non-lensed sources of the same class at comparable distances, and that lensing can be efficiently used to improve the constraints on neutrino emission from FSRQs. The strongest limit is obtained for the *Fermi*-detected B0218+357; it corresponds to a total neutrino power of $1.1 \times 10^{46}\,\mathrm{erg\,s^{-1}}$, about one order of magnitude lower than the lowest limit achieved with non-lensed FSRQs. This limit is expected to improve in the future, in particular when the multi-km^3 neutrino telescope KM3NeT [50], with an instrumented volume about 100 times bigger than that of ANTARES, becomes operational in the Mediterranean.

A similar study could in fact be performed with the lensed quasar QSO 0957+561, a doubly-imaged, wide-separation system with the source at $z = 1.41$ and the lensing galaxy at $z = 0.36$ [11, 12]. This quasar is not in the ANTARES field of view, but its neutrino emission could be constrained by IceCube; the information for this source is therefore also included in table 1. Assuming a typical ANTARES sensitivity for a similar source in the ANTARES field of view, as given in [28], the 90% C.L. upper limit on the neutrino luminosity for QSO 0957+561 would be $L^{90} \simeq 1.8 \times 10^{47}\,\mathrm{erg\,s^{-1}}$. Based on the sensitivities presented in [51],
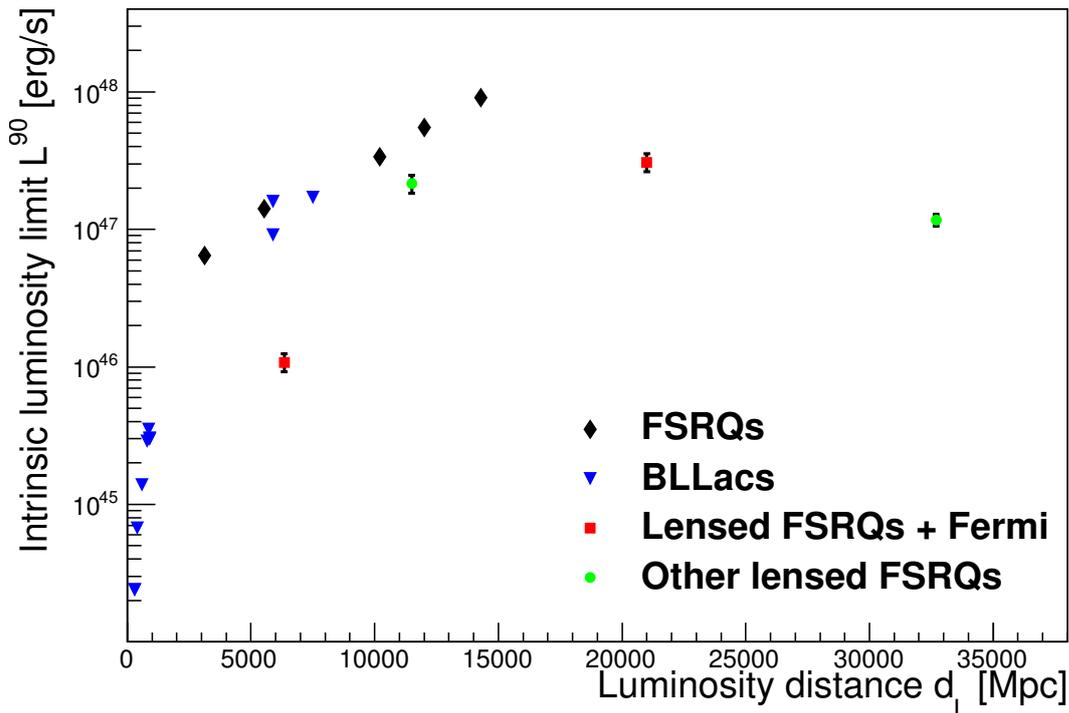
Figure 1. Limit on the instrinsic neutrino luminosity L^{90} of blazars viewed by ANTARES as a function of their luminosity distance. The black diamonds (resp. blue triangles) correspond to non-lensed FSRQs (resp. BLLacs) in the list of candidate sources of [28]. The red squares correspond to the two lensed FSRQs also seen by *Fermi* (B0218+357 and PKS 1830–211) and the green circles correspond to the two other lensed blazars considered in this paper (B1422+231 and B1030+074). The error bars on the lensed sources represent the uncertainty in the determination of the magnification factor.

the limit set by IceCube would approximately be 10 times stronger, and thus comparable to the limit obtained by ANTARES on B0218+357.

In conclusion, we suggest that neutrino telescopes include the lensed FSRQs discussed above in their future searches for steady point-source neutrino emission.

Acknowledgments

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